

In this lecture we will recap the concept of Eddington luminosity, and discuss the lifetime of Pop III stars.

04-21-2016

We have so far figured out; for Pop III stars:

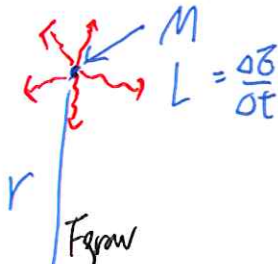
$$M_* \sim 100 M_\odot$$

$$R_* \sim 5 R_\odot$$

$$T_{\text{eff}} \sim 10^5 \text{ K}$$

$$L \sim 10^6 L_\odot \lesssim L_{\text{Edd}}$$

• Eddington luminosity.



σ_T :
cross-sectional
area for the
electron to intercept
a photon.

$$\text{flux} = \frac{\frac{dG}{dt}}{\Delta A} = \frac{L}{4\pi r^2}$$

$$\frac{dG}{dt} = \text{flux} \times \sigma_T = \frac{L}{4\pi r^2} \cdot \sigma_T$$

$$F_{\text{rad}} = \frac{dG}{dt} = \frac{L}{c \cdot dt}$$

Setting $F_{\text{rad}} = F_{\text{grav}}$,

$$L_{\text{Edd}} = \frac{4\pi c G M \cdot m_H}{\sigma_T}$$

OR

$$L_{\text{Edd}} = 10^{38} \text{ erg s}^{-1} \cdot \left(\frac{M}{M_\odot} \right)$$

Then we discussed the luminosity of the LIGO gravitational wave source GW150914, 04-21-2016

Before merging:

36 $M_{\odot}$ 29 $M_{\odot}$



After merging
62 $M_{\odot}$

$$\Delta E = 3 M_{\odot} c^2$$

$$\Delta t \lesssim 1 \text{ s.}$$

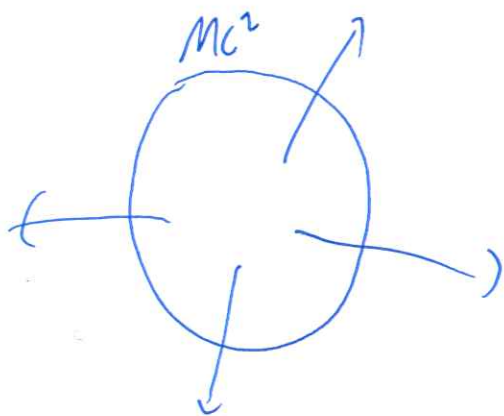
$$L_{\text{GW}} \approx \frac{\Delta E}{\Delta t} \sim \frac{10^{54} \text{ erg}}{1 \text{ s}} = 10^{54} \text{ erg s}^{-1}$$

Compare $L_{\text{add}} = 10^{38} \cdot \left(\frac{M}{M_{\odot}}\right)$

$$\sim 10^{40} \text{ erg s}^{-1} \quad (62 M_{\odot} \sim 100 M_{\odot})$$

$$\Rightarrow L_{\text{GW}} \gg L_{\text{add}} !$$

• Nature's Maximum Luminosity



$$L_{\text{max}} = \frac{Mc^2}{\Delta t_{\text{min}}}$$

$$\Delta t_{\text{min}} = \frac{R}{c} \quad (\text{causality})$$

$$L_{\text{max}} = \frac{Mc^2}{R/c} = \frac{Mc^3}{R_{\text{min}}}$$

$$\rightarrow R_{\text{min}} \geq R_s = \frac{2GM}{c^2} \sim \frac{GM}{c^2}$$

$$\therefore L_{\text{max}} = \frac{Mc^2}{(GM/c^2)} = \frac{c^5}{G} = 10^{59} \text{ erg s}^{-1}$$

$$L_{\text{Einstein}} = c^5/G$$

04-21-2016

— Einstein luminosity

• Lifetime of Pop III star.

$$t_{*, \text{Pop III}} \simeq \frac{\Delta E}{L} \sim \frac{\Delta E}{L_{\text{add}}}$$

• Stars live off nuclear fusion

$$\Delta E_{\text{nuc}} = \eta_{\text{fusion}} \cdot Mc^2$$

$$\eta_{\text{fusion}} = 0.007 = 0.7\%$$

$$\therefore t_{*, \text{Pop III}} \simeq \frac{\Delta E_{\text{nuc}}}{L_{\text{add}}} = \frac{\eta_{\text{fusion}} \cdot Mc^2}{L_{\text{add}}} \sim 3 \times 10^6 \text{ yr}$$

Compare to age of universe at $z=20$

$$t_H(z=20) \sim 10^8 \text{ yr.}$$

$\Rightarrow$ Pop III stars live fast!
