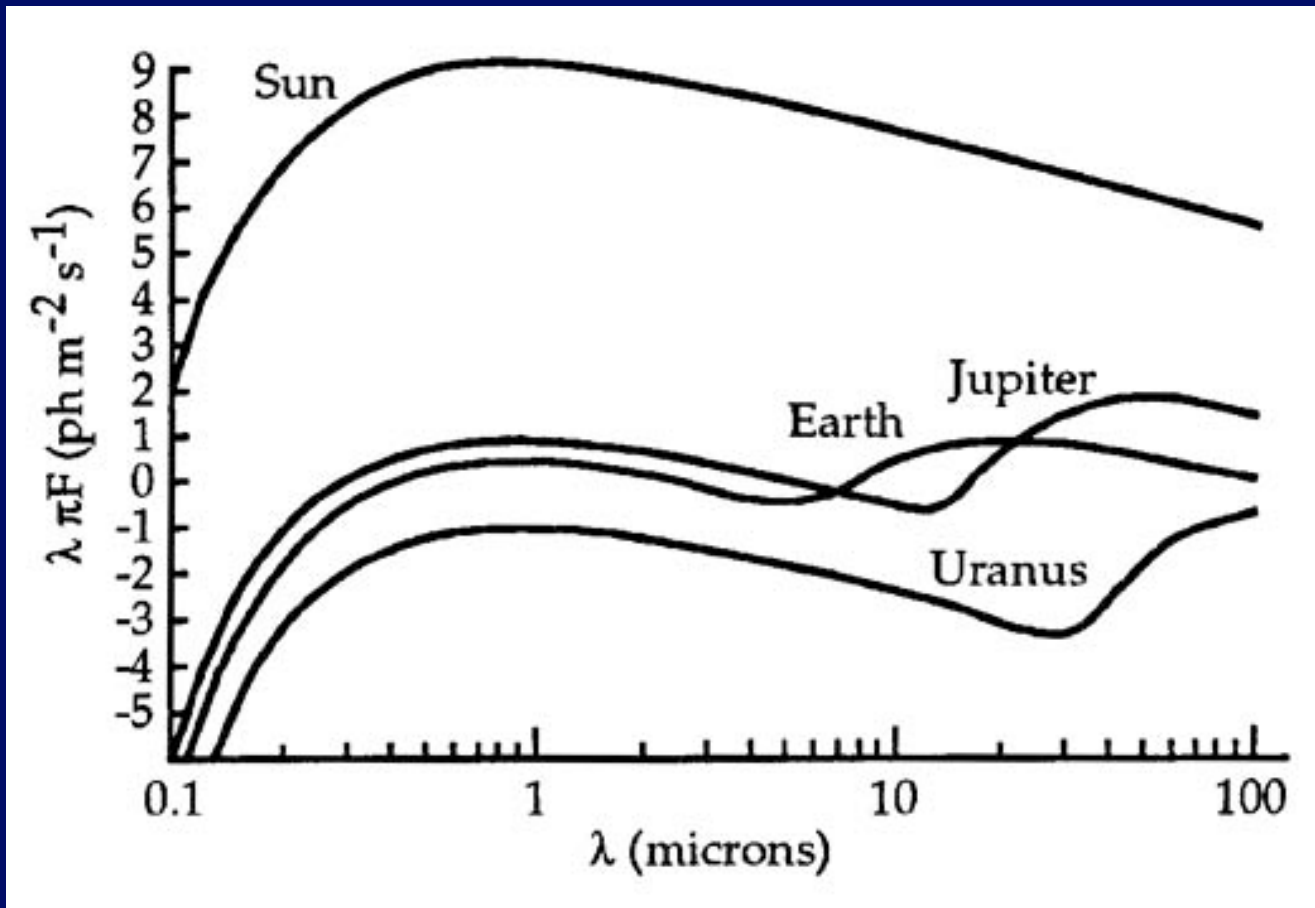
Planet Detection

Estimating f_p

Can We See Them?

- Not easily
  - 4 planets claimed, but planets very far from star, so some doubts
- Problem is separating planet light from star light
  - Star is 10^9 times brighter in visible light
  - “Only” 10^6 times brighter in infrared

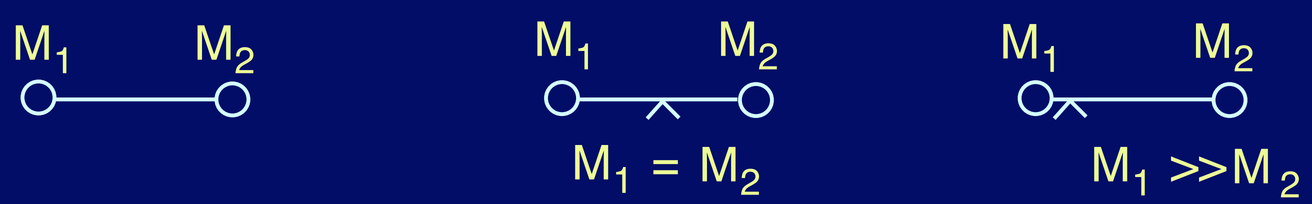
Planet is Much Fainter than Star



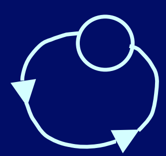
Indirect Detection

Wobbling star

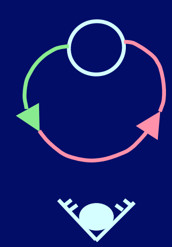
Detect effect of planet on star (both orbit around center of mass)



Large planet will make a star “wobble”



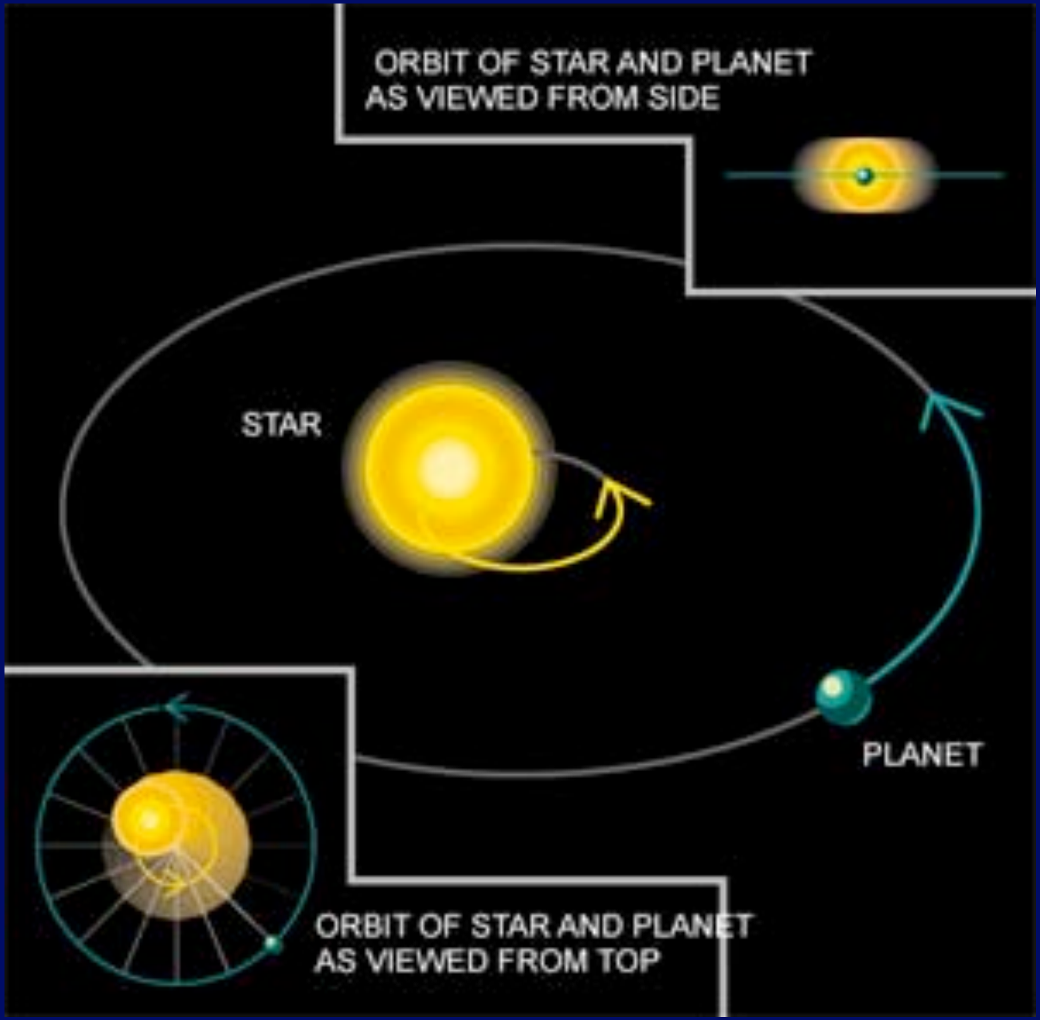
In plane of sky observe position shift



Along our line of sight

Observe Doppler Shift

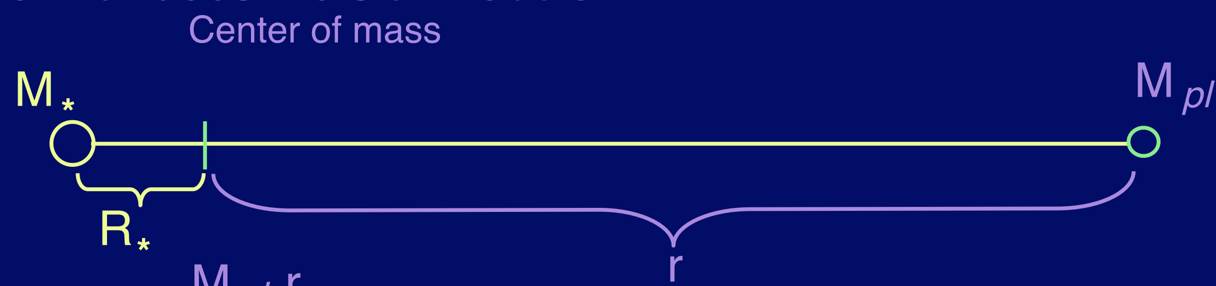
Star and Planet Orbit Center of Mass



The Astrometric Technique

Measure stellar position (angle) accurately - see wobble compared to more distant stars

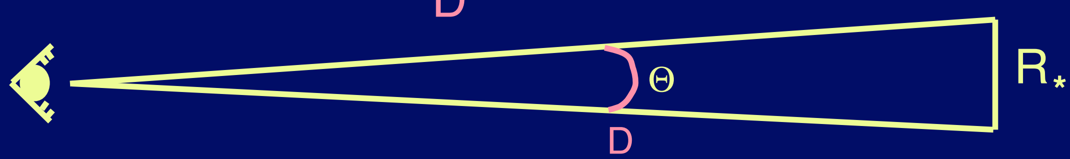
How far does the star wobble?



$$R_* = \frac{M_{pl}\, r}{M_*}$$

We measure angle; for small angles,

$$\Theta = \frac{R_*}{D} \qquad \text{in radians}$$

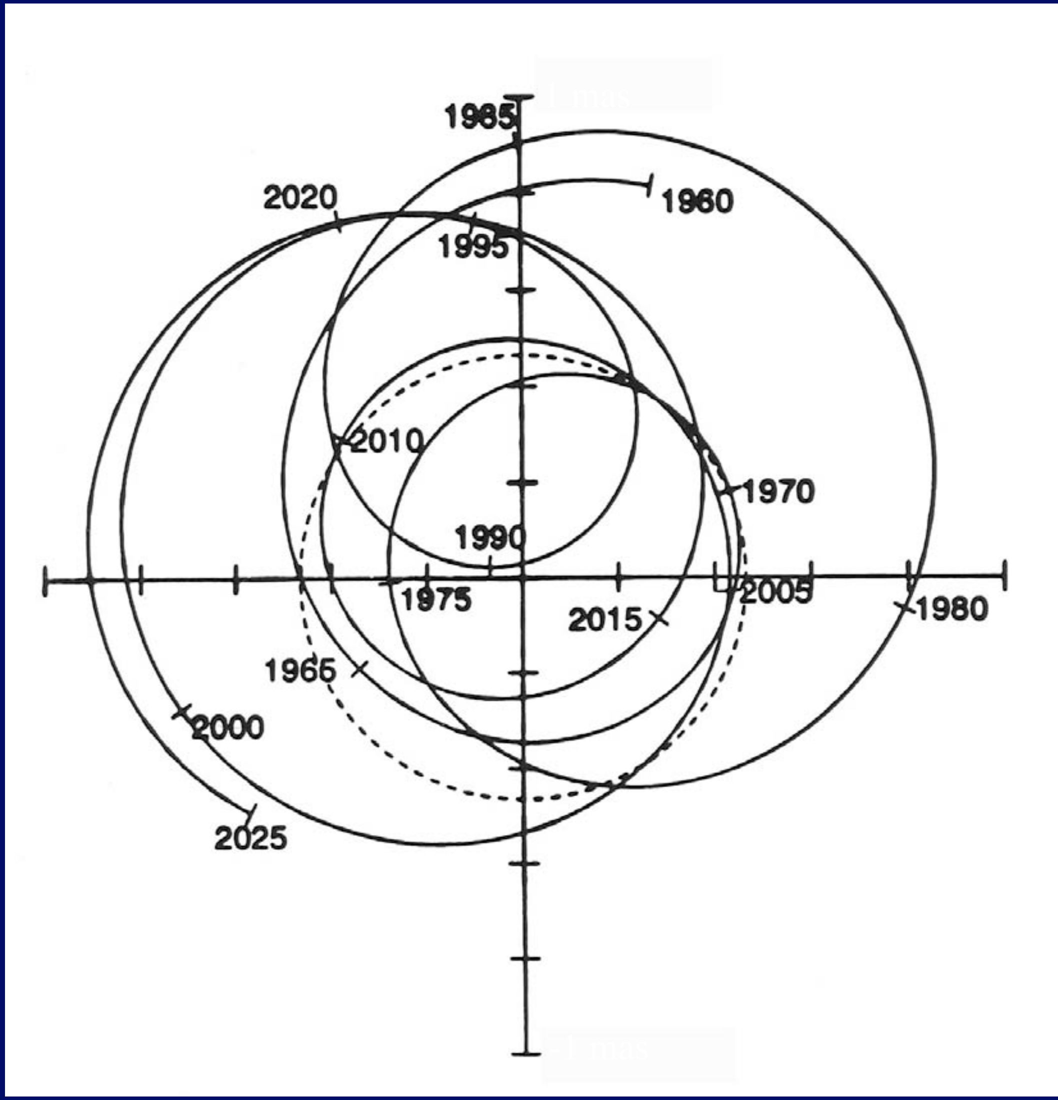


so $$\Theta = \frac{M_{pl}\, r}{M_*} \frac{1}{D}$$

Big planet, big orbit
small star, close to sun

Current limit: 1 mas = 10^{-3} arcsec = 2.8×10^{-6} degrees
= 4.9×10^{-8} radians

e.g. $M_{pl} = M_{Jupiter}$, $M_* = M_\odot$, $D = 15\ ly \Rightarrow \Theta = 1$ mas



The Sun as viewed from 10 pc (~30 *ly*)

30 *ly*

Planet	M_P (M_J)	R (AU)	P (years)	$V_\star$ ($m\,s^{-1}$)	Θ at 10 pc (mas)
Mercury	1.74E-4	0.387	0.241	0.008	6.4E-6
Venus	2.56E-3	0.723	0.615	0.086	1.8E-4
Earth	3.15E-3	1.000	1.000	0.089	3.0E-4
Mars	3.38E-4	1.524	1.881	0.008	4.9E-5
Jupiter	1.0	5.203	11.86	12.4	0.497
Saturn	0.299	9.54	29.46	2.75	0.273
Uranus	0.046	19.18	84.01	0.297	0.084
Neptune	0.054	30.06	164.8	0.281	0.156
Pluto	6.3E-6	39.44	247.7	3E-5	2.4E-5

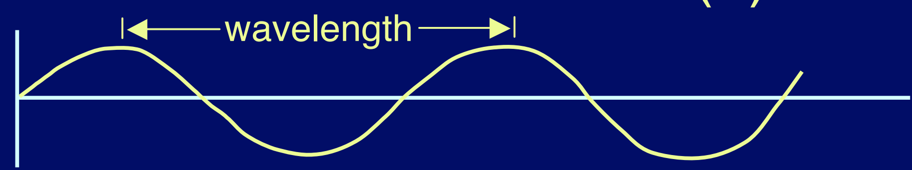
The Spectroscopic Method

- Relies on Doppler Effect
- Motion of star towards and away from us
- Almost all planets around other stars found by this method so far

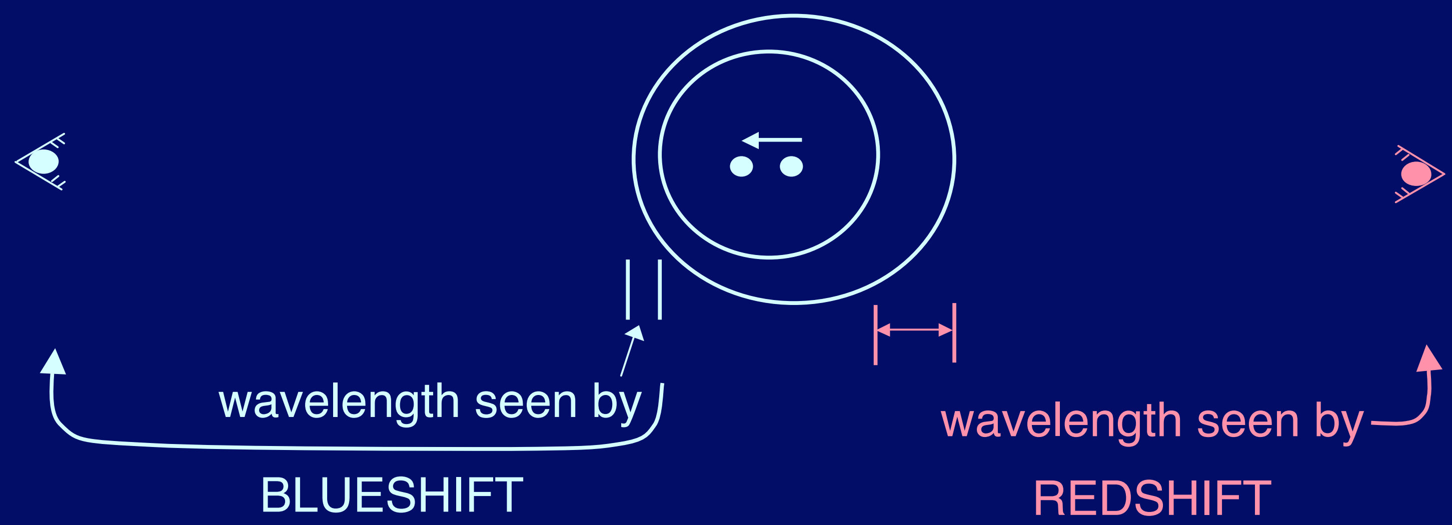
The Doppler Shift

Light is a wave

(λ)



moving star



$$\frac{\lambda_{\text{observed}}}{\lambda_{\text{emitted}}} = 1 + \frac{v}{c}$$

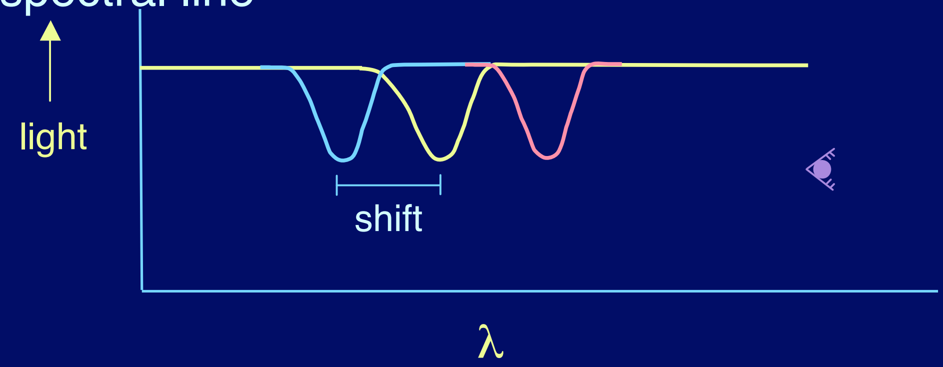
Doppler Shift ⟶ Magnitude and direction of velocity

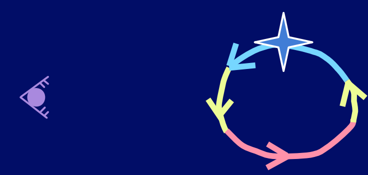
But only along line-of-sight

The Spectroscopic Technique

Measure velocity, not position, of star

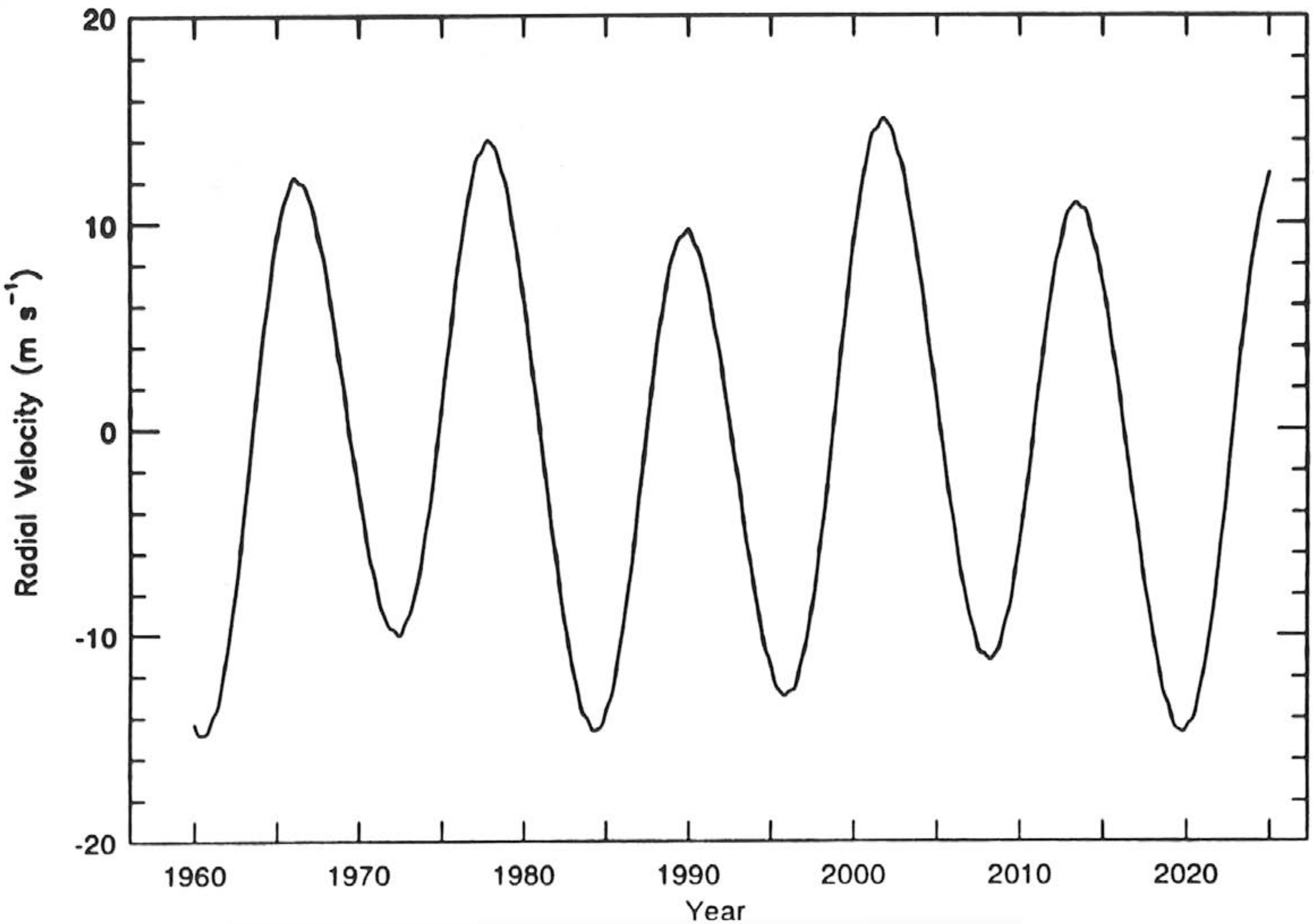
Use spectrometer to get Doppler Shift of spectral line





$$\text{Shift} \propto V_* \propto \frac{M_{pl}}{M_*^{1/2}\, r^{1/2}}$$

Big planet, small orbit
Small star
Distance doesn't matter (except for brightness)
Edge - On



Motion of the Sun caused by Jupiter, ...

30 *ly*

Planet	M_P (M_J)	R (AU)	P (years)	$V_\star$ ($m\,s^{-1}$)	Θ at 10 pc (mas)
Mercury	1.74E-4	0.387	0.241	0.008	6.4E-6
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What We Can Learn

1. There is a planet

   (If not a mistake)

2. The orbital period (P)

   (The time for pattern to repeat)

3. The orbital radius

   $$r^3 \propto M_* P^2$$

   (Kepler's Third Law)

4. Lower limit to planet mass (M_{planet})

   Conservation of momentum ⟶

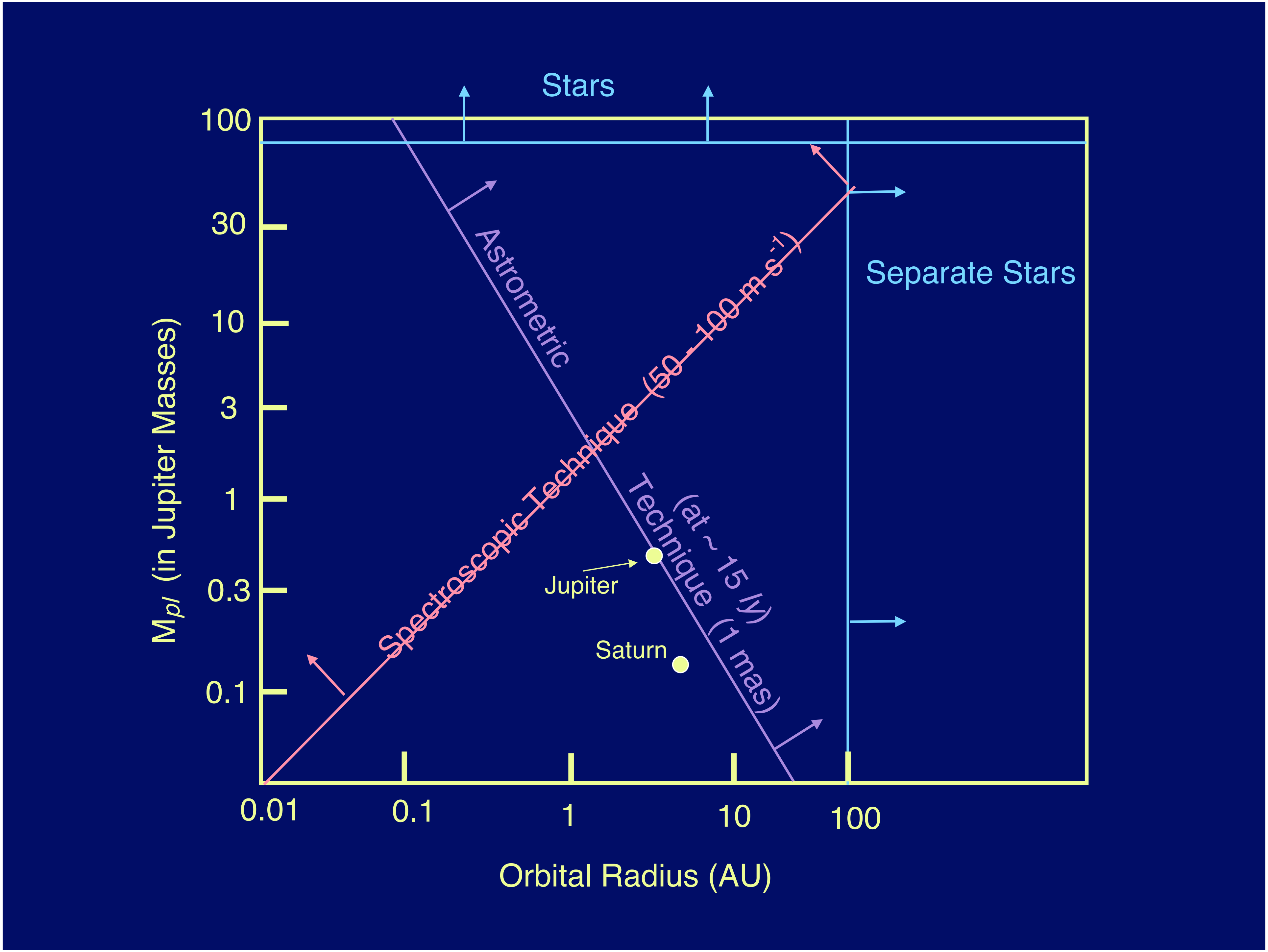
   $$M_{pl} \geqslant \frac{M_* V_* P}{2\pi r}$$

   = if we see orbit edge-on

   > if tilted

Comparison of Search Methods

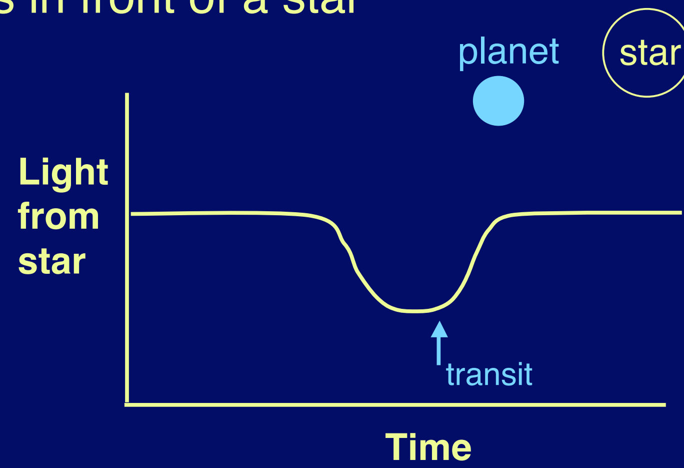
Advantages	Astrometric	Spectroscopic
	Big Planet	Big Planet
	Big Orbit	Small Orbit
	Small Star	Small Star
	Nearby Star	--
		Edge-on Orbit



Stars
Separate Stars
Astrometric Technique (at ~ 15 ly) (1 mas)
Spectroscopic Technique (50 - 100 m s-1)
Jupiter
Saturn
100
30
10
3
1
0.3
0.1
Mpl (in Jupiter Masses)
0.01
0.1
1
10
100
Orbital Radius (AU)

Other Methods

Transits: Planet passes in front of a star

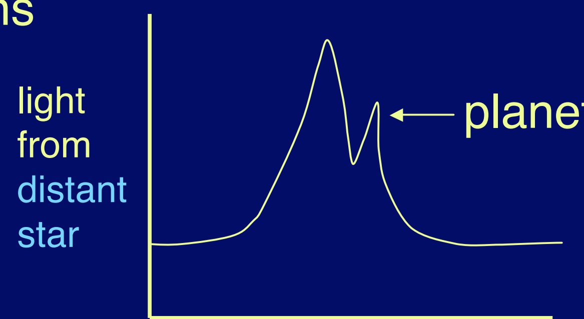


Only about 0.5% of stars with planets will line up

First planet found with this method in January 2003; 35 detected as of January 2008, many more coming soon…

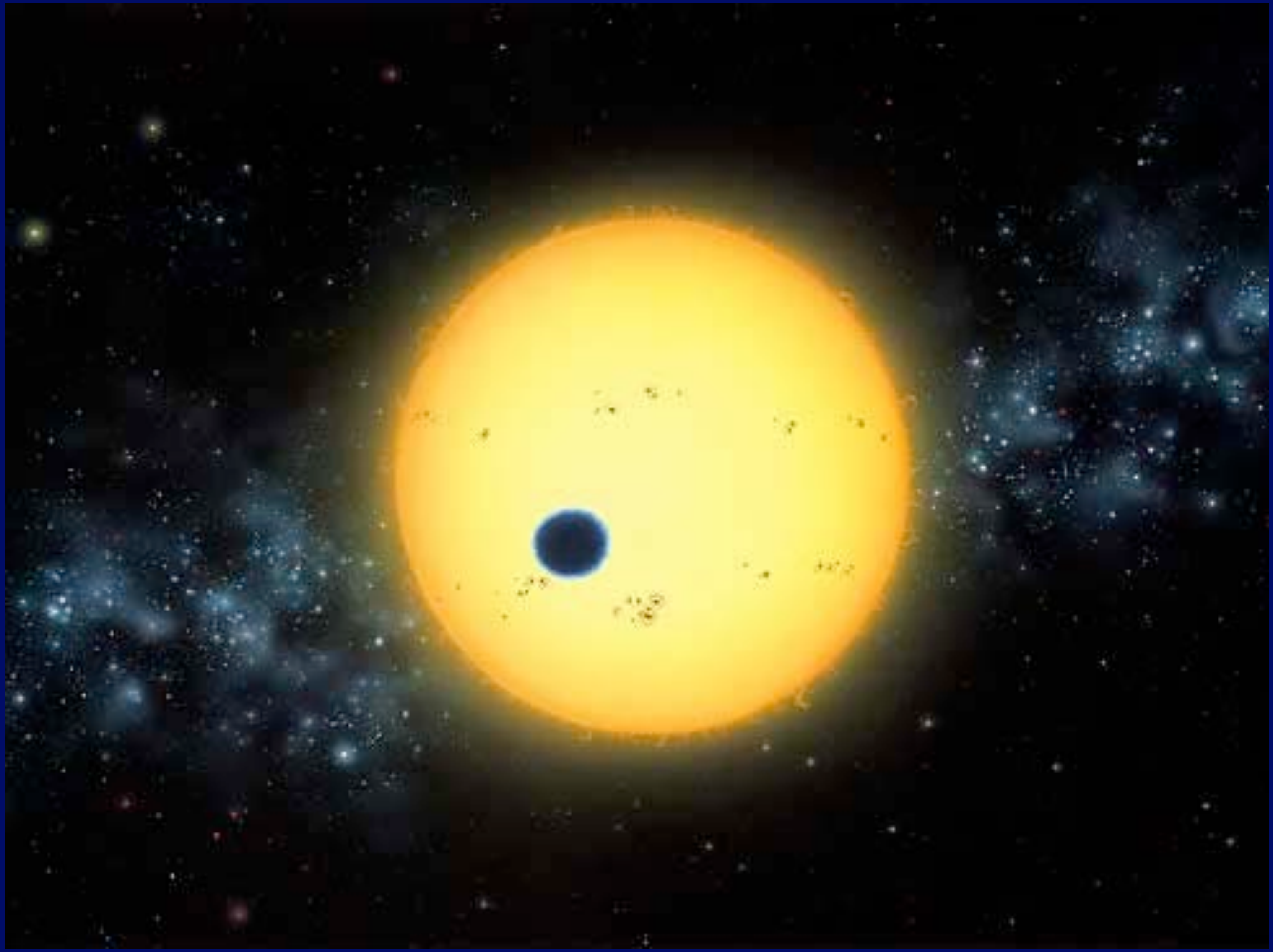
Microlensing: Light from more distant star is focused by gravity of nearer star passing in front

Fortuitous alignment ⇒ brightens



Four planets found this way as of January 2008

Artist's conception of Transit of HD209458

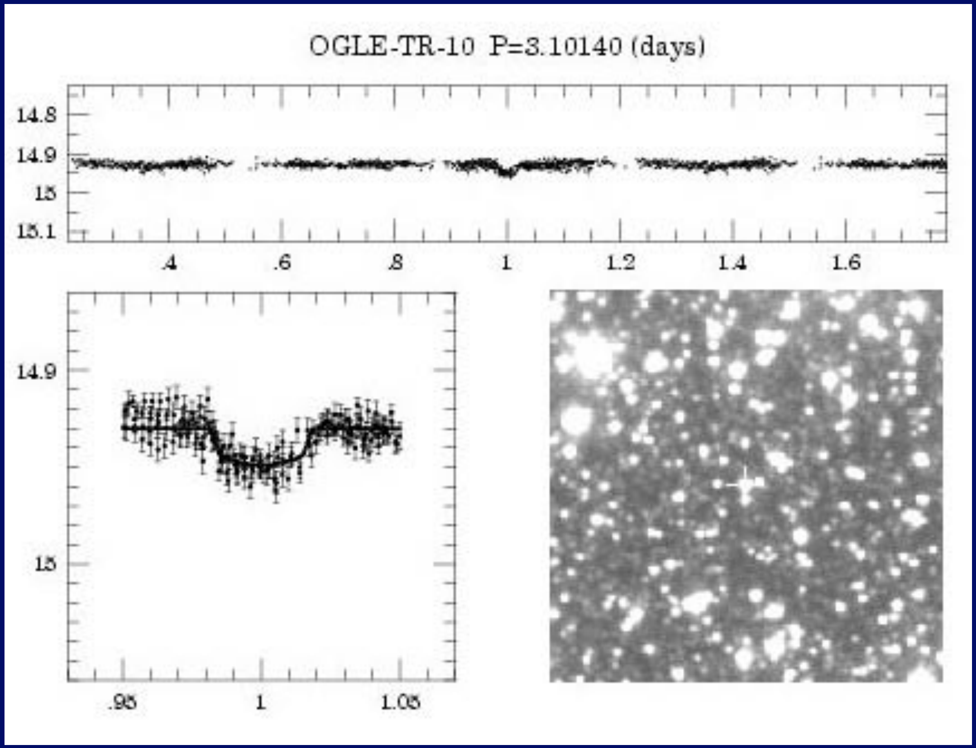


http://www.extrasolar.spaceart.org

Timing

- Delays or advances in periodic signals
  - Pulses from pulsar
    - First planets found that way in 1992
    - Not suitable for life!
  - Oscillations in white dwarfs
    - First found this way in 2007 by grad student at UT

Planets from the Transit Method

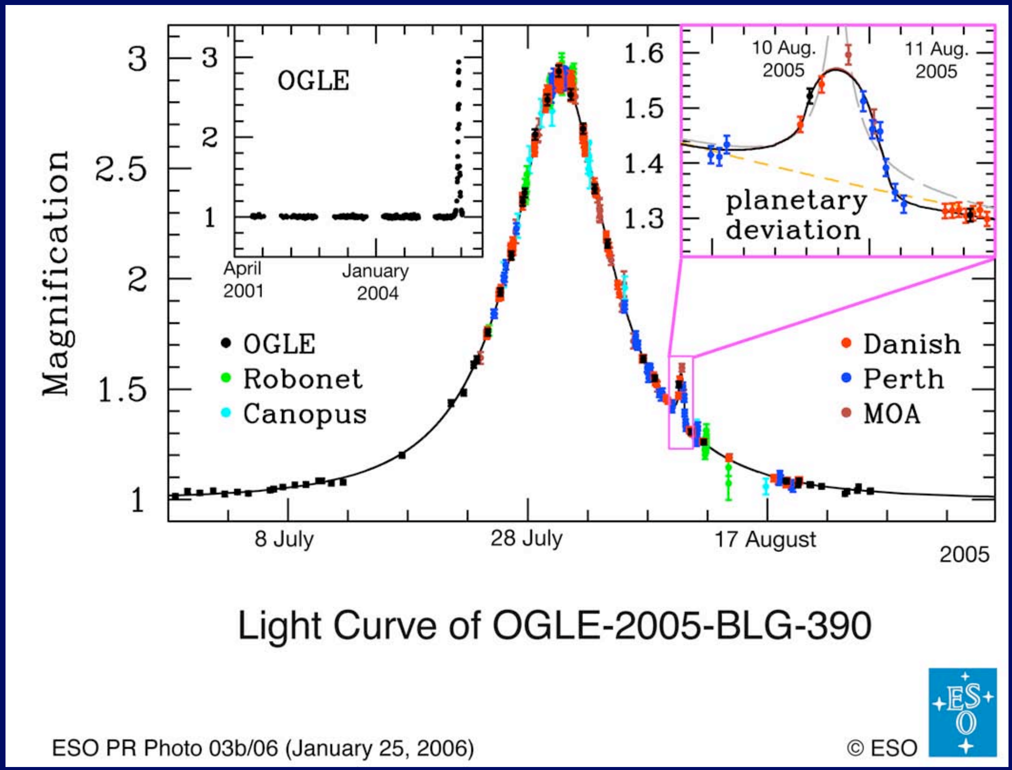


Light curve

Star field, shows star

OGLE-TR-10

Planet Detected by Microlensing



Sharp spike indicates second lens. Mass of second lens only 8 x 10^{-5} as massive as star. Most likely mass of planet is 5.5 M_{earth} and separation from star is 2.6 AU. Most likely star is low mass (0.22 M_{sun}).

This method can detect very low mass planets, but they are one-time events. Cannot follow up.

OGLE 2005-BLG-235Lb, announced 1/25/06

http://www.eso.org/outreach/press-rel/pr-2006/pr-03-06.html

Current Statistics (Jan. 2008)

- Based on Extrasolar Planets Encyclopedia
  - http://exoplanet.eu/
- 271 Planets in 221 systems
- 26 with multiple planets
- Most planets in one system is 4 (55 Cancri)
- Least massive
  - $M = 0.0158\ M_{Jup} = 4.8\ M_{earth}$ (Gliese 581c)

Implications of New Planets

Planets more massive than Jupiter can form around stars like the Sun.

Large Planets can form much closer to a star than Jupiter (or move there)

Does this mean we are unusual and our ideas about other planetary systems are just “solar system chauvinism”?

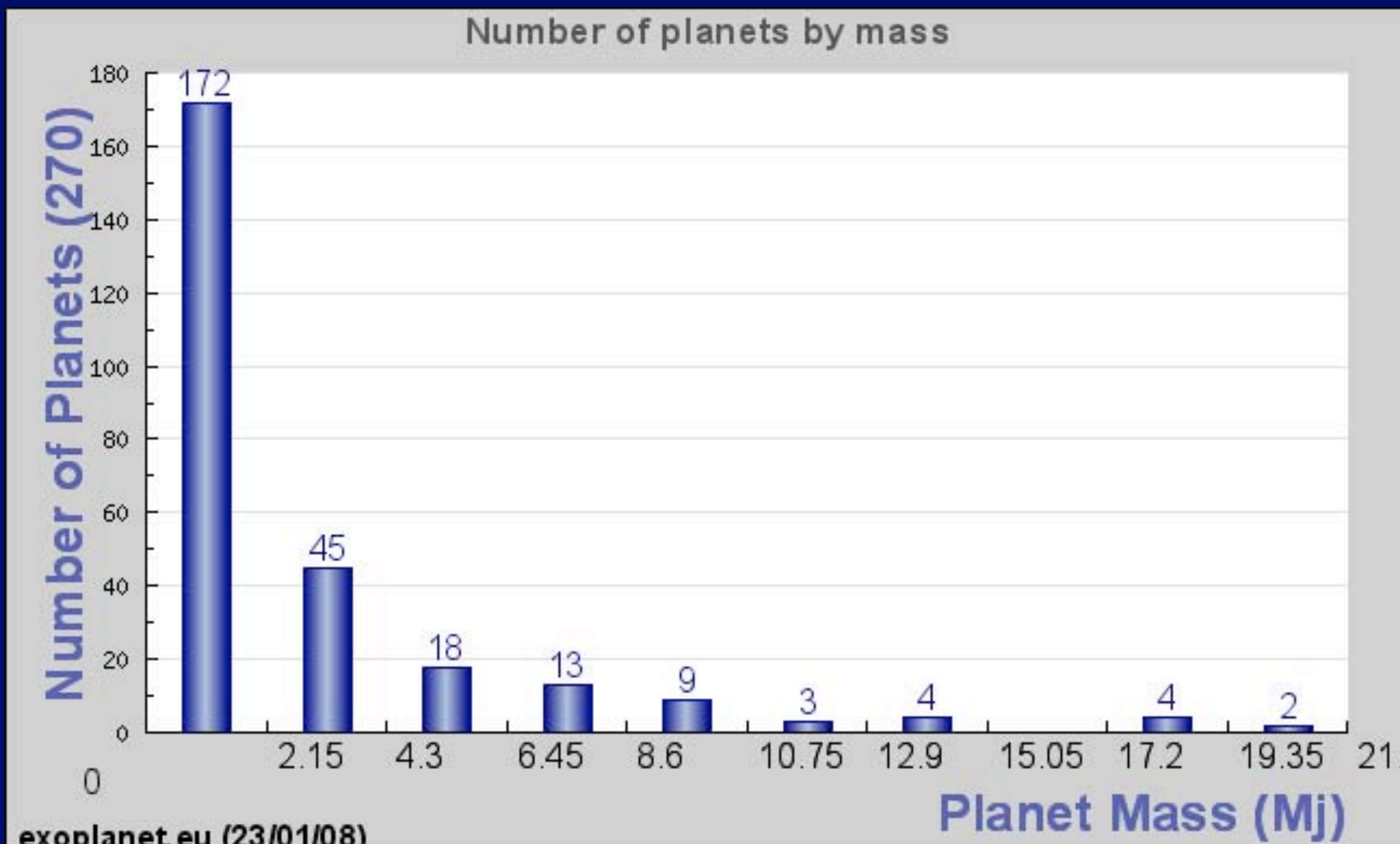
Not necessarily.

The ones found so far are the “easy” ones. (Big planets close to a star)

Now there are many more with lower masses than higher masses.

Too early to say that we are unusual.

Number of planets for different masses

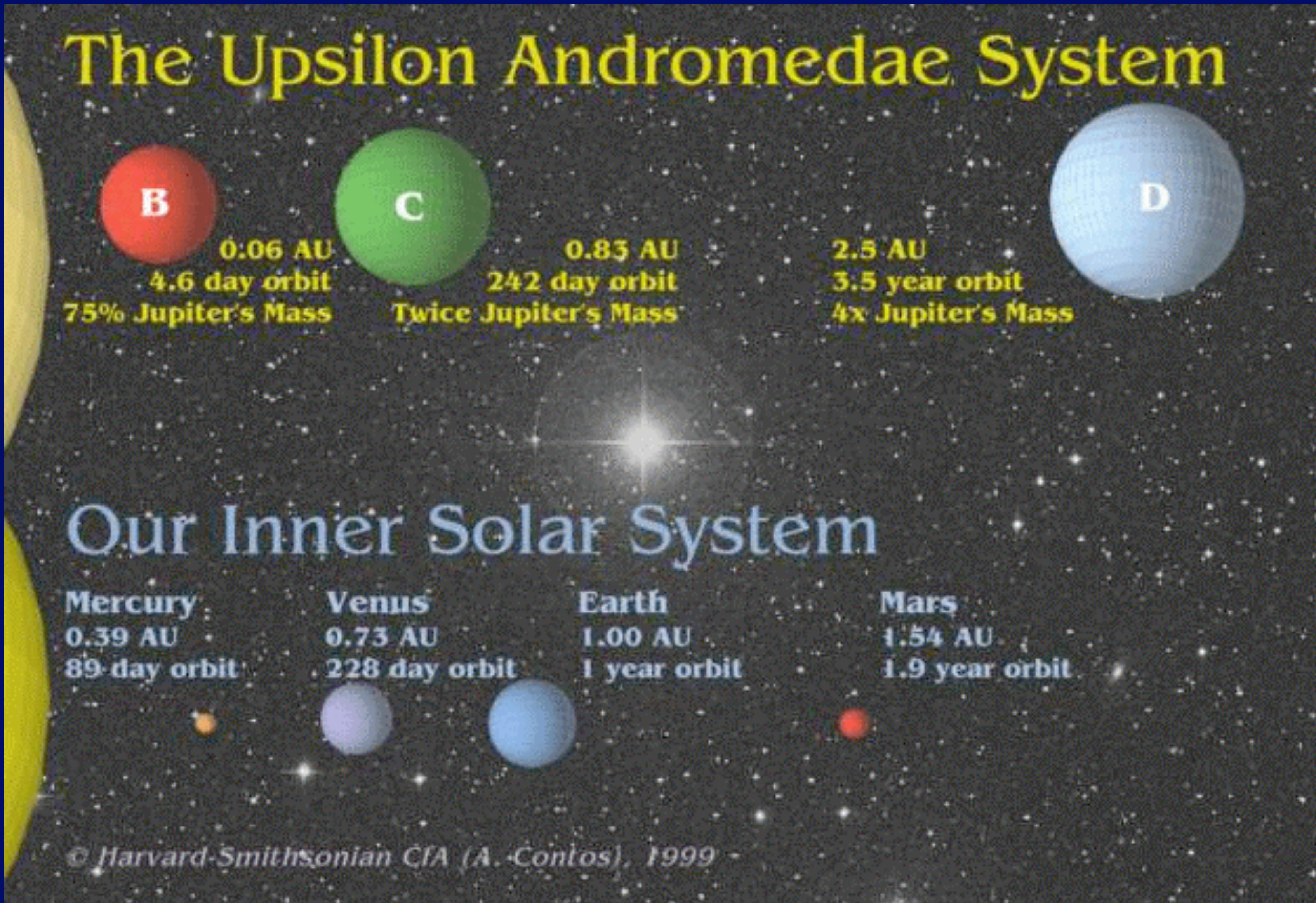


Estimating f_p

- Maximum? $f_p \sim 1$
  - All young stars may have disks
- Binaries?
  - Can have disks, but planet formation?
  - Even if form planets, orbits may not be stable
  - If reject binaries, $f_p < 0.3$

Estimating f_p

- Minimum?
  - Based on success rate of searches ($n_{found}/n_{searched}$)
  - Estimates now up to 5% ($f_p > 0.05$)
    - Note larger than 0.02 given in book
  - Extrapolate trends to finding
    - Smaller planets, larger orbits, …
  - Estimates range from 0.11 to 0.25
- Allowed range: $f_p = 0.05$ to 1.0
  - Explain your choice!
  - Include/exclude binaries?


The Upsilon Andromedae System
B
0.06 AU
4.6 day orbit
75% Jupiter's Mass
C
0.83 AU
242 day orbit
Twice Jupiter's Mass
D
2.5 AU
3.5 year orbit
4x Jupiter's Mass
Our Inner Solar System
Mercury
0.39 AU
89 day orbit
Venus
0.73 AU
228 day orbit
Earth
1.00 AU
1 year orbit
Mars
1.54 AU
1.9 year orbit
© Harvard-Smithsonian CfA (A. Contos), 1999

Future Prospects

Transits

CoRoT Dec. 2006-present

Has reported two planets as of Jan. 2008

Kepler (2009)

Monitor 100,000 stars for 4 years

"Hundreds of Terrestrial Planets"

Astrometric Method

GAIA ~ 2012

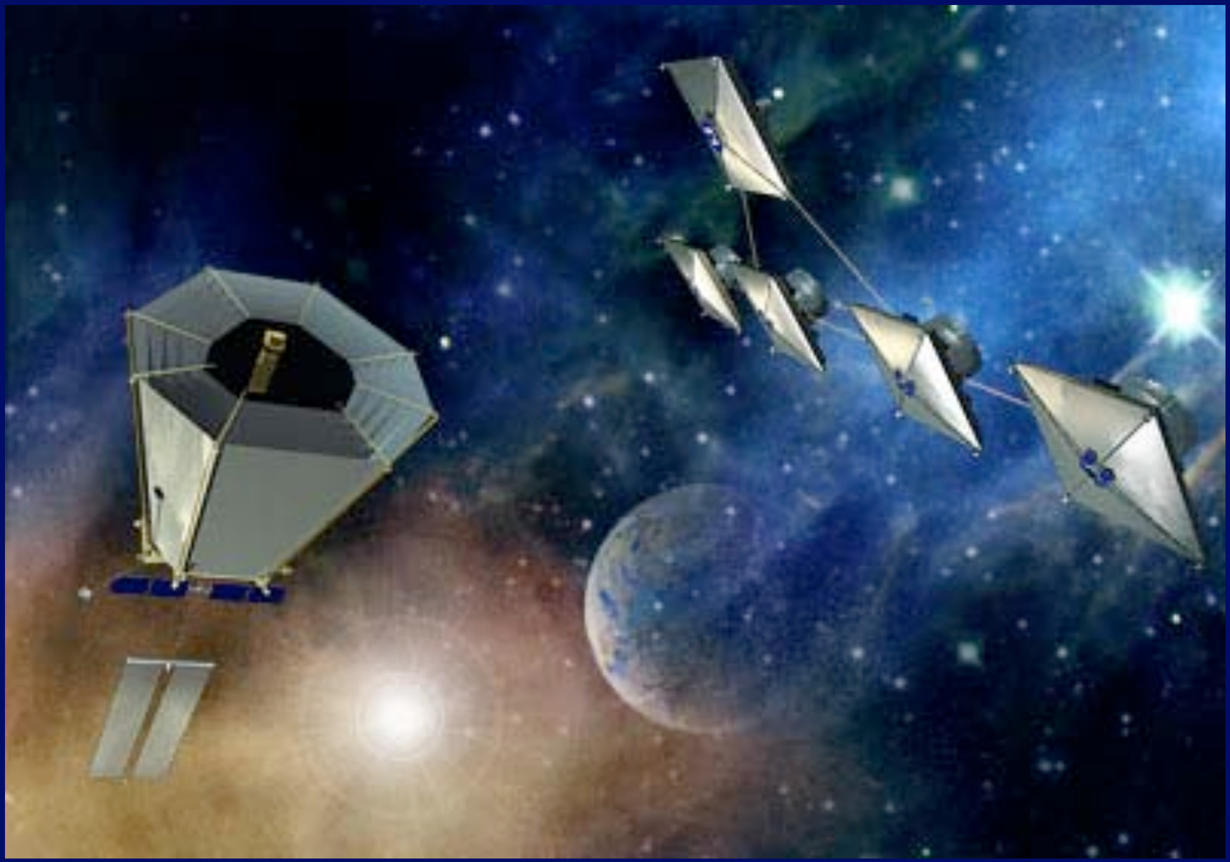
M_J Planets out to 600 ly.

Direct Detection in Future

- Terrestrial Planet Finder (TPF)/Darwin
  - TPF-C Visible light coronagraph (~2014)
  - TPF-I Infrared interferometer (~2020)
- Goal is to detect earth-mass planets
- And to see what gases in atmosphere
  - Suitable for life?
- http://planetquest.jpl.nasa.gov/TPF/tpf_index.html

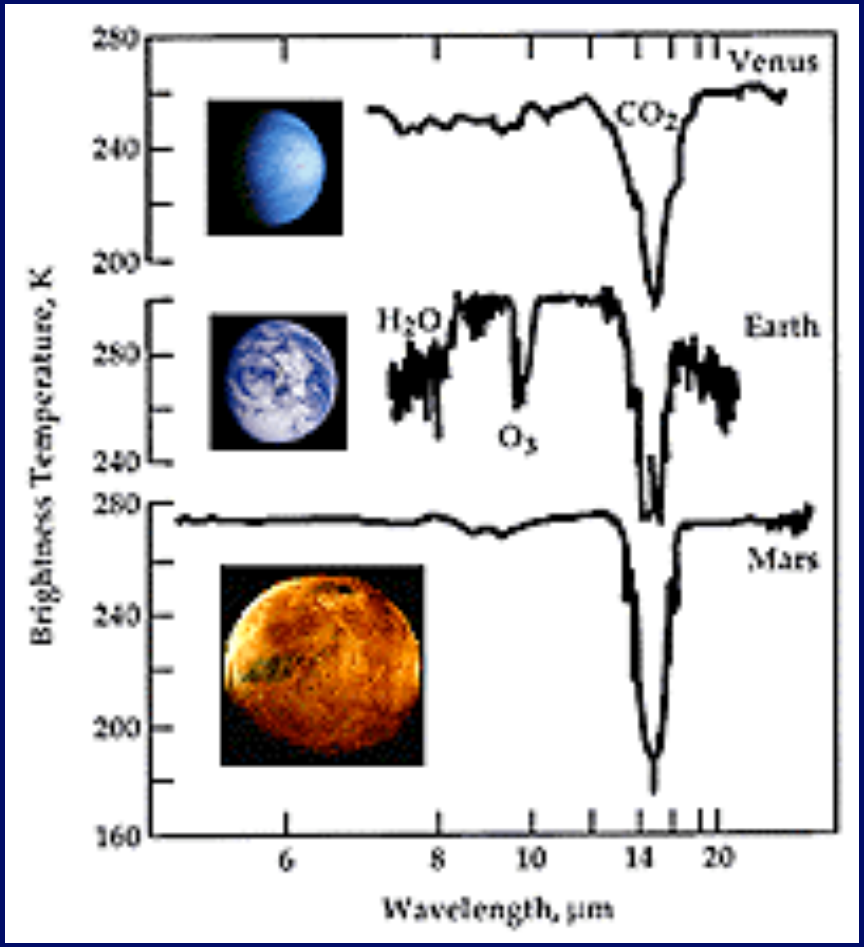
TPF Concepts

TPF-I Infrared Interferometer (2020)



TPF-C Visible light coronagraph (2014)

Spectroscopy of atmosphere



Planet Detection Methods

Michael Perryman, Rep. Prog. Phys., 2000, 63, 1209 (updated November 2004)

[corrections or suggestions please to michael.perryman@esa.int]