Life in the Outer Solar System

Jupiter



Big $R = 11R_{\oplus}$ Massive $M = 300 M_{\oplus}$ $= 2.5 \times all$ the restDay about 10 Earth hoursYear about 12 Earth yearsThick Atmosphere, mostly H₂, HeBut also more complex moleculesColors, storms

Like Miller - Urey

Life in Jupiter Atmosphere?

Sagan-Salpeter, etc.

Sinkers Floaters Hunters (Plankton) (Fish) (Fish)



Galileo Results on Jupiter Reached Jupiter Dec. 1995 Sent probe into Jupiter's atmosphere at 100,000 mile/hour Decelerated at 230 g Lasted for 57 min.

Found: Strong winds Turbulence, little lightning _<

Life less likely?

Surprise: Little or no H₂O May have entered in an unusual place (fewer clouds)





Europa (Moon of Jupiter) Surface: Fractured Ice Subsurface Oceans? (Heated from Inside)



Close-up of "ice floes"



Model of Europa's Interior



Ice crust may be 10-30 km thick. Ocean may be 90 km deep.

Future Missions

- Juno, launched in 2011
 - Will go into polar orbit, map gravity field
 - Determine if Jupiter has a rocky core
- Jupiter Icy Moon Explorer (JUICE)
 - ESA Selected in 2012
 - Launch 2022, arrive 2030
 - Ganymede, Callisto, Europa orbiter
 - Look for evidence of organic molecules

Saturn



- Big (9.4 R⊕)
- Massive (95 M_{\oplus})
- Year 29.5 earth years
- Day 0.43 earth days
- Composition similar to Jupiter



Titan

- Moon of Saturn
- Diameter ~0.4 Earth
- Atmospheric Pressure = 1.5 × Earth
- 85% Nitrogen BUT
- Cold (~90 K)
- Reducing atmosphere
- Haze
- Lab for prebiotic chemistry

The Cassini-Huygens Mission



- Launched 10/13/97
- Arrived Saturn 7/2004
- Cassini studies
 - Saturn
 - Moons
- Huygens
 - Dropped onto Titan
 - Study atmosphere
 - Surface

http://saturn.jpl.nasa.gov/



Huygens Probe



- Released from Cassini
- Slowed by heat shield
- Parachute deployed
- Soft landing
- Sampled gases in atm.
- Results:
 - High winds
 - 430 km/hr at 120 km

Titan Surface 10km up



- Mosaic of images
- Taken during descent
- Clearly shows features

Photo: ESA

Titan



- River channel
- Coastline
- Liquid is present
- Methane (CH₄)

Photo: ESA

Water Rift and Methane Springs?



- Straight feature:
- Water ice extruded?
- Stubby channels:
- Methane springs?

Lakes at northern latitudes

- Radar mapping of northern latitudes (2006)
- Strong evidence for liquid lakes
- And big cloud of ethane (C₂H₆)
- Ethane raining (or snowing) into lakes

Lakes and Islands



Image from Feb. 2007: based on radar. Large lake and island (size of Big Island, Hawaii) And smaller lakes

From the surface of Titan



- First view of surface
- "Rocks" of water ice
 Pebble size (15 cm)
- Surface yielding
- Mixture of ices
 - Water
 - hydrocarbons

More Titan Results

- Hints of ammonia (NH₃)/water (H₂O) ocean
 - About 200 km under surface
 - Outgassing of NH_3 may supply N_2 atm.
- Mapping by radar reveals many lakes and seas of hydrocarbons
 - Seasonal changes in size, depth of a lake
 - Total hydrocarbons on surface about 100 times total oil and gas reserves on Earth

Possible Site for Life

- Miller-Urey type experiments with Titan atm:
 Formed amino acids and nucleotide bases
- Methane-based life?
- Metabolize with H₂ and C₂H₂, produce CH₄
- Parallel to O₂ and glucose, produce CO₂
- Also, could produce atmospheric nitrogen

Enceladus



- Moon of Saturn
- Very shiny
- Part of surface old (craters)
- Part is new, with cracks
- Cassini saw ice geysers (2006)
- Subsurface liquid water
- Source of heat unclear

Geysers on Enceladus



How to search for life

Have to decide what test indicates life Hard to anticipate conditions (recall Viking results) What about finding "protolife"?

National Academy report - how to search for life

- 1. Delivery by comets, meteorites e.g. Mars meteorites
- 2. Sample return Mars possible
- 3. Experiments by landers -Viking on Mars, ...
 Future: Europa probe and return? Titan?
 Issues of contamination
- 4. Biomarkers

Presence of both O_2 and CH_4 in Earth atmosphere indicative of life How convincing?

Detecting Life on Earth from Space Galileo used during close Earth approach Photographs (1 km resolution) No clear signs of intelligent life Spectrometers - evidence of life Lots of O₂



1 a. Galileo long-wavelength-visible and near-infrared spectra of arth over a relatively cloud-free region of the Pacific Ocean, north meo. The incidence and emission angles are 77° and 57° respec-The (b' $\sum_{a}^{e} - X_{a}^{a}$) 0-0 band of 0₂ at 0.76 µm is evident, along a number of H₂0 features. Using several cloud-free regions of 1g airmass, we estimate an 0₂ vertical column density of 1.5 kmjat±25%. b and c, infrared spectra of the Earth in the 2.4-5.2 µm n. The strong v₃ CO₂ band is seen at the 4.3 µm, and water vapour s are found, but not indicated, in the 3.0 µm region. The v₃ band rous oxide, N₂O, is apparent at the edge of the CO₂ band near m, and N₂O combination bands are also seen near 4.0 µm. The



methane (0010) vibrational transition is evident at 3.31 μ m. A cr estimate¹⁰ of the CH₄ and N₂O column abundances is, for both spec of the order of 1 cm-amagate (=1 cm path at STP).

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RADIO EMISSION: CLEAR EVIDENCE OF INTELLIGENT LIFE

Spectroscopy of atmosphere



Could be detected with future large space telescope, but very difficult Need specialized capabilities